Analysis and model-based optimization of a pectin extraction process

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A complex dynamo inferred from the hemispheric dichotomy of Jupiter’s magnetic field

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The Juno spacecraft, which is in a polar orbit around Jupiter, is providing direct measurements of the planet's magnetic field close to its surface1. A recent analysis of observations of Jupiter's magnetic field from eight (of the first nine) Juno orbits has provided a spherical-harmonic reference model (JRM09)2 of Jupiter's magnetic field outside the planet. This model is of particular interest for understanding processes in Jupiter's magnetosphere, but to study the field within the planet and thus the dynamo mechanism that is responsible for generating Jupiter's main magnetic field, alternative models are preferred. Here we report maps of the magnetic field at a range of depths within Jupiter. We find that Jupiter's magnetic field is different from all other known planetary magnetic fields. Within Jupiter, most of the flux emerges from the dynamo region in a narrow band in the northern hemisphere, some of which returns through an intense, isolated flux patch near the equator. Elsewhere, the field is much weaker. The non-dipolar part of the field is confined almost entirely to the northern hemisphere, so there the field is strongly non-dipolar and in the southern hemisphere it is predominantly dipolar. We suggest that Jupiter’s dynamo, unlike Earth’s, does not operate in a thick, homogeneous shell, and we propose that this unexpected field morphology arises from radial variations, possibly including layering, in density or electrical conductivity, or both.

Unlike Earth, for which the top of the dynamo region is well defined by the core–mantle boundary—that is, the boundary between the electrically conducting liquid-iron outer core (in which dynamo action occurs) and the overlying, poorly conducting rocky mantle—for Jupiter the corresponding region is less clearly defined. Even though self-sustaining dynamo action is most probably confined to depths below the metallic-hydrogen transition, the field may be affected by flow in the overlying molecular-hydrogen region3–5, which may have substantial electrical conductivity, especially close to the depth of the metallic-hydrogen transition6. Accordingly, we map the field at four equally spaced radii from the surface of Jupiter (corresponding to \( r = R_J \pm 0.37 R_J \)), at which the electrical conductivity is vanishingly small, to \( r = 0.85 R_J \), the likely depth of the metallic-hydrogen transition.

To do so requires mapping the field below the orbit of the spacecraft, and so we must address the instability due to downward continuation. We do so by regularizing the solution using a quadratic norm based on the horizontal Laplacian of the radial magnetic field, thereby finding the smoothest possible map of the field for a given fit to the observations8. We select Juno magnetometer observations1 from eight orbits in the radial distance range from \( r = 1.06 R_J \) (periopause) to \( r = 2.2 R_J \) (roughly corresponding to Juno’s highest latitude), take 30-s averages of the data (corresponding to one rotation of the spacecraft) and weight the data according to an estimate of their measurement uncertainty. Our resulting dataset consists of 1,991 observations of each of the three components of the magnetic field.

In Fig. 1 we show maps of the radial component of the magnetic field at a range of depths using our regularized inversion from the surface to \( r = 0.85 R_J \) and compare with JRM092. At all depths, positive radial flux in the northern hemisphere is confined to a band (the northern-hemisphere flux band), which becomes narrower with depth. Some of the flux from this band then re-enters through an intense spot on the equator7 (the Great Blue Spot), at a longitude of around 90° west (in System III coordinates). The morphology of the magnetic field lines is shown in Fig. 2 (an animated version of Fig. 2 is available at https://doi.org/10.6084/m9.figshare.6828953). Elsewhere, and corresponding to a large proportion of the surface, the radial flux is much weaker.

The narrowing of the northern-hemisphere flux band with depth, and more generally the concentration of flux into increasingly localized regions with depth rather than, for example, the emergence of more small-scale spots, is surprising given our intuition acquired from mapping Earth’s magnetic field at depth. It suggests that Jupiter’s magnetic field at depth may be morphologically simpler than expected. This field morphology and its contrast to Earth’s field is particularly apparent in Fig. 3, in which we show the non-dipolar part of the field (at \( r = 0.90 R_J \)) and, for comparison, Earth’s non-dipole field (at Earth’s core–mantle boundary). Jupiter’s non-dipole field is almost entirely confined to the northern hemisphere, where the non-dipole field peaks at 3 mT, a value almost three times stronger than the peak dipolar field. Jupiter’s field is dipolar in the southern hemisphere and largely non-dipolar in the northern hemisphere, unlike Earth’s field.

The strong concentration of magnetic flux in the northern-hemisphere flux band and in the Great Blue Spot implies the existence of large horizontal magnetic field gradients at the borders of these features, which would suggest that strong secular (temporal) variation of the magnetic field is likely. For example, around the Great Blue Spot the gradient in the radial field is approximately 3 mT/(106 m); with an assumed flow speed of the order of 10−4 m s−1 (the lower end of estimates of flow speed10,11), we might therefore expect secular variation of the order of 103 nT yr−1. Although high, this estimate is not necessarily inconsistent with earlier inferences of much weaker time dependency12 because secular variation at such small spatial scales would be strongly attenuated at the altitude of the previous observations. In addition, this estimate will be reduced if the flow is preferentially orthogonal to the field gradient, although for the Great Blue Spot that is unlikely on geometrical grounds to be the case. Therefore, we believe that the Great Blue Spot offers a very promising opportunity for forthcoming Juno orbits to detect secular variation.

Numerical dynamo models in simple homogeneous shells typically produce fields that are either strongly dipolar or dominated by multipolar fields10,13. Jupiter’s field is neither, being predominantly dipolar in one hemisphere and non-dipolar in the other, suggesting that the field is not generated in a simple homogeneous region. Here we consider several possible explanations. First, we consider the possibility, although unlikely, that we have observed the field in a rare transitional
Jupiter's core will be soluble in hydrogen at the temperature and pressure expected there\textsuperscript{23}. This may lead to gradual core dissolution, and may have been crucial in Jupiter's thermal history\textsuperscript{26,27}. Dissolution of rock and ice in metallic hydrogen will increase the density of the hydrogen region. Recent Juno observations of Jupiter's gravity field are consistent with the existence of a partially or fully dissolved core inside Jupiter, with rock and ice non-uniformly mixed in the hydrogen out to approximately half the radius of the planet\textsuperscript{25}; the region further out may be homogeneous, except for helium rain.

If, as theory and observations suggest, the metallic-hydrogen region is layered (the upper layer solute-free and the lower layer containing dissolved rock and ice), the implications for the dynamo will depend on the convective instability of these layers. The upper layer is most probably convectively unstable, given the very large heat flux observed at Jupiter. The properties of the lower layer are far less clear. If the lower layer is stable, then dynamo action will be confined to the upper layer and will therefore operate in a shell with a radius ratio (inner to outer radii) of approximately 0.5. A similar geometry has been investigated previously as a possible explanation for the magnetic fields of Uranus and Neptune\textsuperscript{29}, albeit with a numerical dynamo model much less sophisticated than what is now feasible. The magnetic field map obtained from this simulation with a radius ratio of 0.5 (see figure 16, model 5 in ref. 29) bears similarity to the map of Jupiter's field shown here, but with an axial dipole that is much less dominant. In addition, structure may arise from double diffusive convection\textsuperscript{28}.

Alternatively, if the lower layer is convectively unstable, then it could be convecting separately from the layer above owing to the possible presence of a density jump at the boundary between the layers\textsuperscript{28}. Convection in Jupiter's metallic-hydrogen region can be driven by relative density variations ($\Delta \rho / \rho$) of the order of $10^{-6}$, so even a small density jump could be impervious to convection. In this scenario, dynamo action may occur separately in the thick lower shell.
Fig. 3 | Non-dipole field. a, The non-dipolar part of Jupiter’s radial magnetic field at \( r = 0.90R_J \). b, For comparison, the non-dipolar part of Earth’s radial magnetic field at the core–mantle boundary \( (r = 0.55R_E = 3,485 \text{ km}, \text{where} \ R_E \text{ is Earth’s radius}) \). Almost all of Jupiter’s non-dipole radial field is concentrated in the northern hemisphere, whereas Earth’s field is evenly distributed throughout.

(radius ratio of less than 0.2) and in the thin upper shell (radius ratio of approximately 0.5), with the resultant field sharing properties of both a thick-shell dynamo (strong axial dipole) and a relatively thin-shell dynamo (hemispheric asymmetry).

The presence or absence of reduced magnetic flux at high latitude may provide a means of distinguishing between these alternatives. If the lower layer is stably stratified, then convection in the outer layer within the tangent cylinder (the axial cylinder tangential to the interface between the two layers) may differ from that outside the tangent cylinder. If the lower layer is convectively unstable, then such an effect seems less likely to occur. To resolve this additional Juno orbits are required. Juno’s orbit, with perijove precessing northward by approximately 1° per orbit, is evolving in such a way that mid- and high-latitude structure will be better resolved towards the second half of the planned 34-orbit baseline mission.

Data availability

The Juno magnetometer data used in this study will be made available through the NASA Planetary Data System (https://pds.nasa.gov) in accordance with NASA policy. An animated version of Fig. 2 is available at https://doi.org/10.6084/m9.figshare.6828953.

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